

NASA/CR-1998-208710
ICASE Report No. 98-40



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National Aeronautics and
Space Administration

Langley Research Center
Hampton, Virginia 23681-2199

Prepared for Langley Research Center
under Contract NAS1-97046

September 1998

EFFECT OF FINITE COMPUTATIONAL DOMAIN ON TURBULENCE SCALING LAW IN BOTH PHYSICAL AND SPECTRAL SPACES

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Abstract. The well-known translation between the power law of energy spectrum and that of the correlation function or the second order structure function has been widely used in analyzing random data. Here, we show that the translation is valid only in proper scaling regimes. The regimes of valid translation are different for the correlation function and the structure function. Indeed, they do not overlap. Furthermore, in practice, the power laws exist only for a finite range of scales. We show that this finite range makes the translation inexact even in the proper scaling regime. The error depends on the scaling exponent. The current findings are applicable to data analysis in fluid turbulence and other stochastic systems.

Key words. direct numerical simulations, scaling laws, finite range effects

Subject classification. Fluid Mechanics

1. Introduction. A fundamental question in turbulence research is the scaling laws of various physical quantities in the so-called inertial range, where those quantities usually scale as power laws. The scaling exponents are the focus of study. A famous example is Kolmogorov's $-5/3$ law for turbulence energy in the inertial range. This scaling behavior can be observed both in the physical space from the structure functions and in the spectral space from the energy spectrum. It is a common practice to relate power laws in the spectral space with those in the physical space as follows: given a homogeneous random field with a energy spectrum $E(k) \sim k^{-n}$, its correlation function, $C(r)$, scales as $\sim r^{n-1}$, where k is the wave number in the spectral space and r is distance in the physical space. Similar relation holds for the second order structure function: $S(r) \sim r^{n-1}$.

Mathematically, the above translations hold only when $E(k)$ follows a pure power law that extends to $k = 0$ and ∞ and with proper scaling exponents. These restrictions are not met by many physical problems, including turbulence. In particular, the power law of $E(k)$ may exist only in a finite "inertial range": $k_0 \leq k \leq k_1$, where k_0 and k_1 are the large scale cutoff and the dissipation scale cutoff, respectively. Since the wave number k corresponds to a characteristic length scale $l = 2\pi/k$, the corresponding physical scales in the inertial range is given by $[r_1, r_0]$, where $r_i = 2\pi/k_i$ ($i = 0, 1$). A long scaling range, namely $k_1 \gg k_0$ or $r_0 \gg r_1$, is usually required for the study of the physics in that range. This is achieved, e.g., by high Reynolds number turbulent flows. In this case, the general practice is to ignore the effect of the finite scaling range.

In this paper, we show that the finite scaling range does have important effects. More specifically, we study in detail the translation of power laws from spectral space to physical space in 3-D. We find that due to the finite power law range of $E(k)$, $C(r)$ and $S(r)$ only approximately scale as a power law for $r_1 \ll r \ll r_0$.

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Thus the power law range in the physical space is much shorter than that of the corresponding power law given in the spectral space. This phenomenon was mentioned in Frisch [1], but it has not been carefully analyzed. We show how the error of the power law approximation to $C(r)$ and $S(r)$ depends on the cut-off wave numbers, k_0 and k_1 , as well as the scaling exponent n . The latter is of crucial importance. In fact, if n is outside the proper ranges, one may obtain scaling exponents of $C(r)$ and $S(r)$ independent of n .

2. Pure Power Law. Let $u(\mathbf{x})$ be a random scalar field. Its Fourier representation is

$$u(\mathbf{x}) = \int_{R^d} \hat{u}(\mathbf{k}) e^{i\mathbf{k}\cdot\mathbf{x}} d\mathbf{k}.$$

The correlation function and second order structure function are defined as: $C(\mathbf{x}, \mathbf{x}') = \langle u(\mathbf{x})u(\mathbf{x}') \rangle$; and $S(\mathbf{x}, \mathbf{x}') = \langle |u(\mathbf{x}) - u(\mathbf{x}')|^2 \rangle$. where $\langle \rangle$ denotes an ensemble average. Assuming homogeneity: $\langle \hat{u}(\mathbf{k})\hat{u}(\mathbf{k}') \rangle = Q(\mathbf{k})\delta(\mathbf{k} + \mathbf{k}')$, we have that C and S are functions of $\mathbf{r} = \mathbf{x} - \mathbf{x}'$:

$$(1) \quad C(\mathbf{r}) = \int_{R^d} Q(\mathbf{k}) e^{i\mathbf{k}\cdot\mathbf{r}} d\mathbf{k}; \quad S(\mathbf{r}) = 2 \int_{R^d} Q(\mathbf{k}) (1 - e^{i\mathbf{k}\cdot\mathbf{r}}) d\mathbf{k}.$$

If we further assume $Q(\mathbf{k}) = Q(k)$ ($k = |\mathbf{k}|$), then C and S are functions of $r = |\mathbf{r}|$. The energy spectrum $E(k)$ is given by $E(k) = S_d Q(k) k^{d-1}$, where $S_d = 1, 2\pi$, and 4π in one, two, and three dimensions, respectively. In 3-D, we have

$$(2) \quad C(r) = \int_0^\infty E(k) \frac{\sin kr}{kr} dk; \quad S(r) = 2 \int_0^\infty E(k) \left(1 - \frac{\sin kr}{kr}\right) dk.$$

Similar expressions can be obtained in general d -dimensional space, see e.g. Reed, Lee, and Truong [2]. In the following, we use Eq. 2 to demonstrate the relation between the power law scalings in the spectral and physical spaces. They can also be obtained for $d \neq 3$ in a similar fashion. It should be noted that the assumptions used in deriving Eq.2 are sufficient but not necessary. All derivations below use only Eq.2. Thus, the results are applicable to any random field that satisfy Eq.2. In particular, it can be a vector field.

When $E(k) = k^{-n}$ for $0 < k < \infty$, we have a pure power law in the spectral space. We now derive the corresponding power law in the physical space. Letting $\rho = kr$, from (2) we have $C(r) = C_0 r^{n-1}$, where C_0 is given by an improper integral

$$(3) \quad C_0 = \int_0^\infty \rho^{-(n+1)} \sin \rho d\rho$$

which exists for $-1 < n < 1$. Note that $n < 1$ and $n > -1$ ensure the convergence of C_0 at $\rho = 0$ and $\rho = \infty$, respectively. The value of C_0 can be found in Gradshteyn and Ryzhik [3]:

$$C_0 = \pi/2 \quad n = 0$$

$$(4) \quad C_0 = \begin{cases} \Gamma(-n) \sin(-n\pi/2) & 0 < n < 1, \\ \Gamma(1-n) \cos((1-n)\pi) & -1 < n < 0. \end{cases}$$

Similarly, we have $S(r) = S_0 r^{n-1}$ with

$$(5) \quad S_0 = 2 \int_0^\infty \rho^{-n} \left(1 - \frac{\sin \rho}{\rho}\right) d\rho.$$

Unlike the case for $C(r)$, the improper integral S_0 exists only for $1 < n < 3$. We note that range of n in which C_0 and S_0 exist are different and do not overlap. Therefore, in order to obtain the power law in the physical space one should choose the correlation function or the structure function according to the scaling exponent n .

3. Truncated Power Law. In practice, $E(k)$ does not appear as a pure power law; it scales as a power law only in certain range of k . Thus the integrals in Eq. 3 and Eq.5 contain no singularities and convergence is not a problem. However, we see below that the scaling behaviors of $C(r)$ and $S(r)$ are still dictated by the above convergence conditions for Eq. 3 and Eq. 5, respectively.

In this section we assume $E(k) = k^{-n}$ in the interval $[k_0, k_1]$ ($k_1 \gg k_0$) and $E(k) \equiv 0$ outside the interval. But the results presented in this paper should be valid also if $E(k)$ has a proper cutoff outside the interval. Then in general $C(r) = A(r)r^{n-1}$ and $S(r) = B(r)r^{n-1}$, where

$$(6) \quad A(r) = \int_{k_0 r}^{k_1 r} \rho^{-(n+1)} \sin \rho \, d\rho, B(r) = 2 \int_{k_0 r}^{k_1 r} \rho^{-n} \left(1 - \frac{\sin \rho}{\rho}\right) d\rho.$$

The fact that A and B are functions of r indicates that $C(r)$ and $S(r)$ no longer follow a single power law. On the other hand, if $k_0 r \ll 1$ and $k_1 r \gg 1$, then $A \approx C_0$ and $B \approx S_0$, provided that $n \in (-1, 1)$ and $n \in (1, 3)$, respectively. In this case, $C(r)$ and $S(r)$ are approximately power laws with exponent $n - 1$.

In the following, we analyze the effect of finite inertial range of $E(k)$ and that of n on the scaling of correlation and structure functions. The main idea is to obtain asymptotic expansions for $C(r)$ and $S(r)$ in terms of $k_0 r (\ll 1)$ and $k_1 r (\gg 1)$. We would like to stress that r need not be very small to achieve the expansions. In fact, r is strictly in the physical space inertial range, i.e., $r_1 \ll r \ll r_0$.

To fix the notation, throughout the paper, A_i and α_i denote *generic* constants. These constants depend on n and the dimension of space but independent of r . In addition, A_i are independent of k_0 and k_1 .

3.1. A useful convergent expansion. First let us provide a result which will be frequently used below. We consider the expansion of

$$\int_{k_0 r}^{k_1 r} \rho^{-\mu} \sin \rho \, d\rho = C_0 - \int_0^{k_0 r} \rho^{-\mu} \sin \rho \, d\rho - \int_{k_1 r}^{\infty} \rho^{-\mu} \sin \rho \, d\rho$$

where $0 < \mu < 2$. Denote the first and second integrals on the r.h.s. by I_1 and I_2 . Using the Taylor expansion of $\sin \rho$ for I_1 and integration by parts for I_2 , we obtain

$$I_1 = \frac{(k_0 r)^{2-\mu}}{2-\mu} + O((k_0 r)^{4-\mu}),$$

$$I_2 = (k_1 r)^{-\mu} \cos k_1 r + O((k_1 r)^{-(1+\mu)}),$$

respectively. Therefore, to the leading orders we have

$$(7) \quad \int_{k_0 r}^{k_1 r} \rho^{-\mu} \sin \rho \, d\rho \approx C_0 - \frac{(k_0 r)^{2-\mu}}{2-\mu} - (k_1 r)^{-\mu} \cos(k_1 r).$$

This expansion converges to C_0 as $k_0 r \rightarrow 0$ and $k_1 r \rightarrow \infty$. The convergence becomes slow when μ is close to 2 or 0.

3.2. Power law approximation for correlation function. In the following, we expand $A(r)$ and $B(r)$. When n is in the proper ranges, Eq. 7 gives the desired result. Otherwise, we may repeatedly apply integration by parts to $A(r)$ and $B(r)$ until the exponent of ρ in the remaining integral falls into the range $(0, 2)$. By Eq. 7, this remaining integral gives the constant term in $A(r)$ and $B(r)$, hence the $A_0 r^{n-1}$ term in $C(r)$ and $S(r)$. However, this term may be dominated by the terms generated from the integration by parts, which are functions of r . In these cases, the $A_0 r^{n-1}$ term is retained in the expansions for the purpose of comparison while the other two terms in Eq. 7 are neglected. It should be noted that the repeated

integration by parts just mentioned is always possible if n is not an integer. For simplicity we assume this is true in the expansions below.

Case $-1 < n < 1$.

Since $0 < 1 + n < 2$, it follows immediately from Eq.7 that

$$(8) \quad C(r) \approx r^{n-1} [C_0 + A_1(k_0 r)^{1-n} - \cos(k_1 r)(k_1 r)^{-(n+1)}].$$

Therefore, the power law of $C(r)$ is only an approximate one. In practice, one obtains the scaling exponent of $C(r)$ by fitting it with a power function

$$f_1(r) = \alpha_0 r^m,$$

where α_0 and m are constants to be determined by a least square procedure. Eq. 8 indicates that in order to obtain accurate exponent one should choose $C(r)$ in an r range such that $(k_0 r)^{1-n} \ll 1$ and $(k_1 r)^{-(n+1)} \ll 1$. Note that the proper choice is determined by the truncation modes as well as the exponent. On the other hand, because the leading order error due to the truncation at k_0 cancels with r^{n-1} , we may better fit $C(r)$ using a function of the form

$$(9) \quad f_2(r) = \alpha_0 r^m + \alpha_1.$$

It should be note that the least square fit now involves solving a non-linear system of equations.

Case $n > 1$.

By Taylor expansions of $\sin \rho$ and integration by parts, we obtain from Eq. 7

$$(10) \quad C(r) \approx A_0 r^{n-1} + A_1 k_0^{1-n} + A_2 k_0^{3-n} r^2.$$

This expansion indicates that the error mainly comes from the low-mode truncation and the error *diverges* as $k_0 \rightarrow 0$. Note that the last term is relatively small if $n < 3$. In this case, we can use Eq. 9 to extract the scaling exponent from $C(r)$. But if $n > 3$, the last term dominates the first term since $A_2(k_0 r)^{3-n} \gg A_0$. Thus, *using Eq. 9 one may obtain $k \approx 2$ regardless of the value of n* . A simple fix seems to be including the r^2 term in the fitting function, i.e.,

$$(11) \quad f_3(r) = \alpha_0 r^m + \alpha_1 + \alpha_2 r^2.$$

However, this does not solve the whole problem because the terms neglected in the expansion Eq. 10 may dominate the first term if n is large. For example, if $n > 5$, a $k_0^{5-n} r^4$ term in the expansion becomes larger than the first and hence Eq. 11 does not work.

Case $n < -1$.

We have

$$(12) \quad C(r) \approx A_0 r^{n-1} + A_1 k_1^{-(n+1)} r^{-2} \cos(k_1 r).$$

Noting that for fixed r , $C(r) \rightarrow \infty$ as $k_1 \rightarrow \infty$ due to the second term. This term, while containing fast oscillations, $\cos(k_1 r)$, decays slower than r^{n-1} . Thus, it is dominating. We may replace r^2 in Eq. 11 by r^{-2} to reflect the asymptotics; however, this modification is not helpful because of the rapid oscillations in the coefficient.

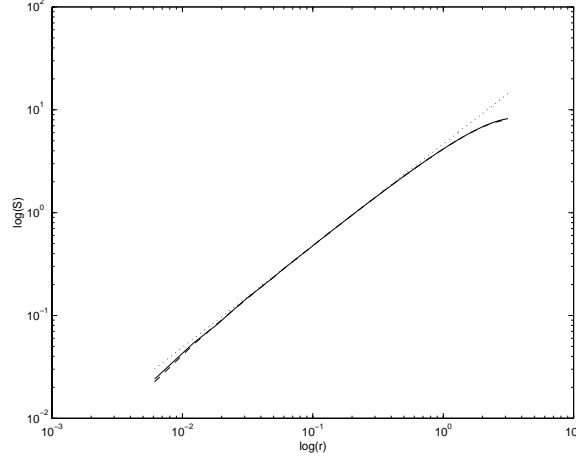


FIG. 1. Log-log plot of $S(r)$ for $n = 2$. Solid line: $S(r)$; dash line: fitted curve using $f_3(r)$ with $m = 1$; dotted line: fitted curve using $f_1(r)$.

3.3. Power law approximation for structure function. The structure function $S(r)$ can be analyzed using the same approach as above. In particular, we observe

$$\begin{aligned}
 & \int \rho^{-n} \left(1 - \frac{\sin \rho}{\rho}\right) d\rho = \int \rho^{-(1+n)} (\rho - \sin \rho) d\rho \\
 (13) \quad & = -\frac{\rho - \sin \rho}{n\rho^n} - \frac{1 - \cos \rho}{n(n-1)\rho^{n-1}} - \frac{1}{n(n-1)} \int \rho^{1-n} \sin \rho d\rho.
 \end{aligned}$$

Thus, the results for $C(r)$ can be applied directly to the last integral.

Case $1 < n < 3$.

It is straightforward to derive

$$(14) \quad S(r) \approx r^{n-1} [A_0 + A_1(k_0 r)^{3-n} + A_2(k_1 r)^{1-n} + A_3(k_1 r)^{-n} \sin(k_1 r)]$$

Note that the last term can be neglected because $k_1 r \gg 1$. The terms with A_1 and A_2 are small if n is not close to 3 and 1, respectively. In this case, one can directly extract the exponent by fitting $S(r)$ with $f_1(r)$. However, $f_3(r)$ is a better fitting function which includes the effect of both terms. This is demonstrated in Fig. 1, where $S(r)$ of a 2-D random field with $E(k) = k^{-2}$ ($k \in [1, 512]$) is plotted. FFT is used in computing $S(r)$ and generating the random field in the $(2\pi)^2$ domain on a 1024^2 lattice. The dotted line is obtained by using f_1 . It has a slope of 0.99, quite close to the exact value, 1. When using f_3 , to avoid solving the nonlinear system of equations, we let $m = 1$ and compute α_i . We have $\alpha_0 = 4.92$, $\alpha_1 = -0.00751$, and $\alpha_2 = -0.757$, indicating that r^{n-1} scaling dominates. In addition, our tests show that the fitting results are sensitive to the data used. The result reported here is obtained from fitting $S(r)$ in the r interval $[0.098, 0.196]$, which satisfies the condition $r_1 \ll r \ll r_0$. Violating this condition renders inaccurate fitted scaling exponent. Furthermore, we note that the Kolmogorov law for homogeneous turbulence, having $n = 5/3$, belongs to this case.

Case $n > 3$.

We have

$$(15) \quad S(r) \approx A_0 r^{n-1} + A_1 k_0^{3-n} r^2.$$

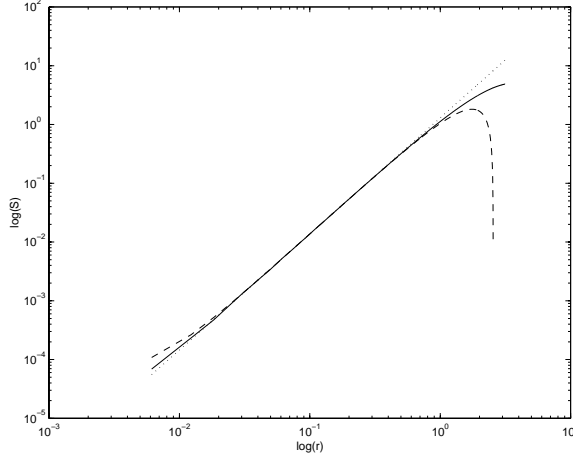


FIG. 2. Log-log plot of $S(r)$ for $n = 4.5$. Solid line: $S(r)$; dash line: fitted curve using $f_3(r)$ with $m = 3.5$; dotted line: fitted curve using $f_1(r)$.

Eq. 15 differs from Eq. 10 only by a constant term. The above test is performed again here with $n = 4.5$. The results are shown in Fig. 2. The slope of the dotted line (fitting with f_1) is 1.98, close to 2 instead of $n - 1 = 3.5$ as predicated by Eq. 15. Letting $m = 3.5$ in f_3 , we find $\alpha_0 = -0.338$, $\alpha_1 = 5.69 \times 10^{-5}$, and $\alpha_2 = 1.38$. Thus r^2 term is indeed larger. Our tests show that increasing n makes α_0/α_2 smaller and hence more dominant r^2 scaling in $S(r)$. It is also seen from the figure that f_3 does not fit better than f_1 due to large n .

Case $n < 1$.

The leading orders of $S(r)$ are

$$(16) \quad S(r) \approx A_0 r^{n-1} + A_1 k_1^{1-n} + A_2 k_1^{-n} r^{-1} \sin(k_1 r).$$

Note that for fixed r , $S(r) \rightarrow \infty$ as $k_1 \rightarrow \infty$. When $n > 0$, Eq. 9 can be used since the last term is small. However, when $n < 0$, we have a situation similar to that of $C(r)$ when $n < -1$.

4. Conclusions. In this paper, we have studied the translation between the power law of energy spectrum and that of the correlation function or the second order structure function. We have obtained the following four conclusions:

1. Power laws in spectral and physical spaces have simple correspondence only for n in proper ranges, i.e., $(-1, 1)$ and $(1, 3)$ respectively for the correlation and structure functions. The effect of finite power law range in the spectral space results in a much shorter power law range in the physical space.
2. Based on the asymptotic expansions, a fitting function f_3 is proposed for better recovery of n from $S(r)$ and $C(r)$. For a given value of n , we recommend using either $S(r)$ ($1 < n < 3$) or $C(r)$ ($-1 < n < 1$) for studying the corresponding power law in the physical space.
3. When n is outside the proper ranges, the correlation and structure functions are generally dominated by some functions of r that are independent of n . Thus recovering n is very difficult, if not impossible.
4. Following the analysis outlined above, one finds that the translation of a power law from the physical space to the spectral space suffers similar problems.

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